

# **Estimating the Depth of Shower Maximum using the Surface Detectors of the Pierre Auger Observatory**

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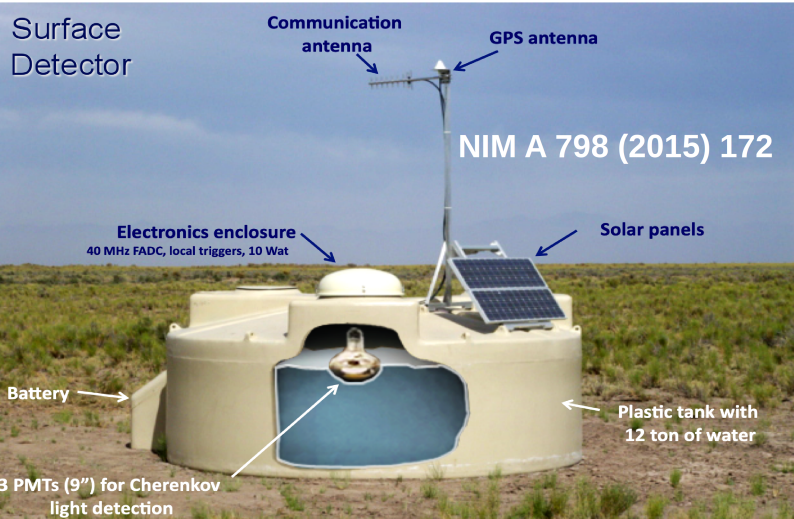
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**36th International Cosmic Ray Conference**

# Surface Detectors of Auger Observatory



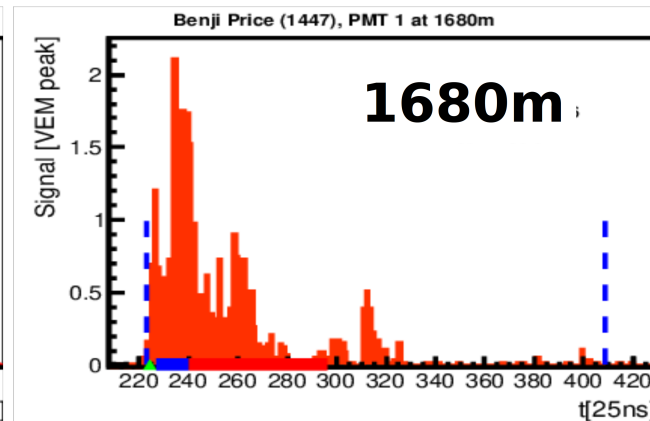
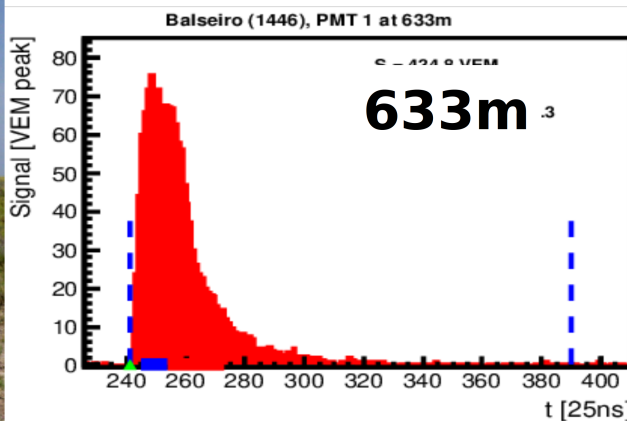
## Water-Cherenkov arrays

- 1500 m spacing; 1600 detectors; (3000 km<sup>2</sup>)
- 750 m spacing; 60 detectors; (23.5 km<sup>2</sup>)

Record a mix of electromagnetic and muonic signals with the muonic signals dominant at large distances

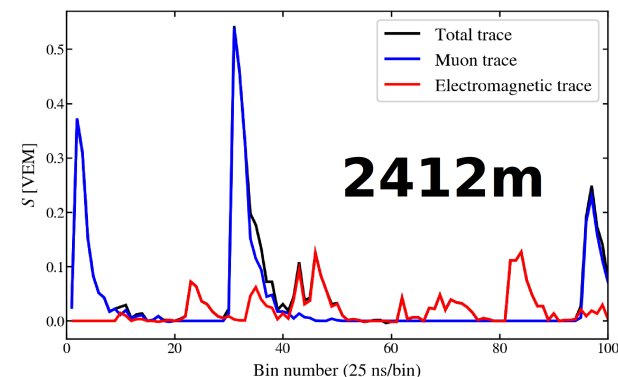
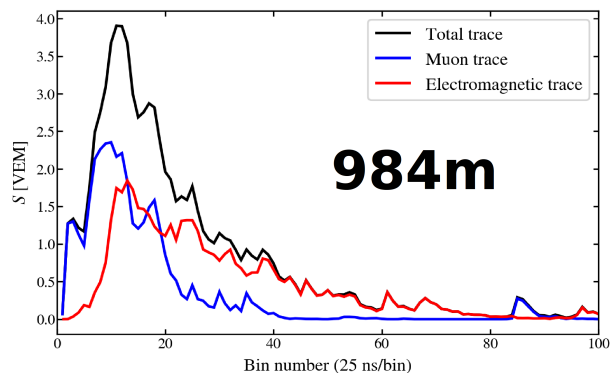
## Data

$E = 17.3 \text{ EeV}$ ,  $\theta = 32^\circ$



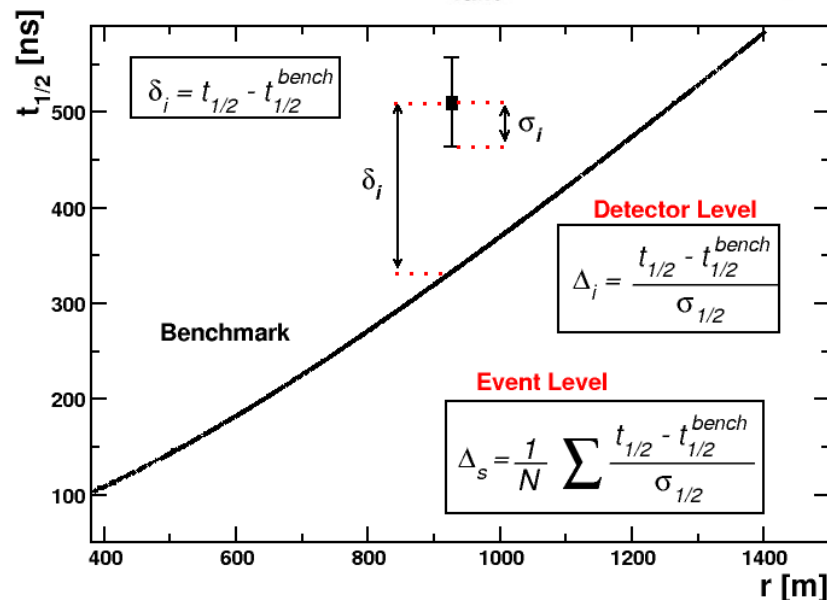
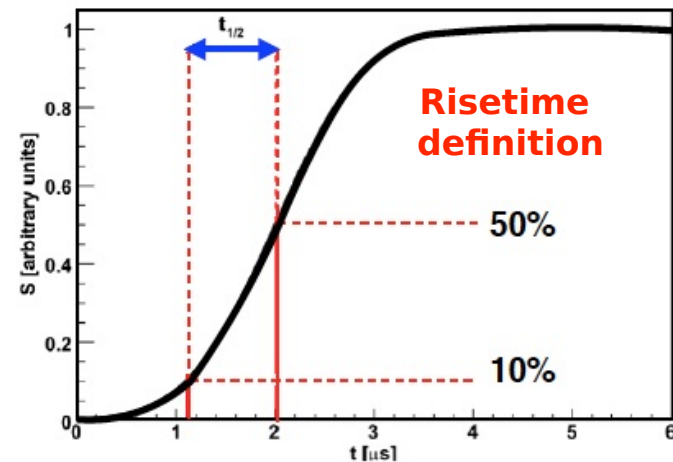
## Simulations

$E = 19.4 \text{ EeV}$ ,  $\theta = 43^\circ$



# The $\Delta$ -method in a nutshell

- Measure risetime ( $t_{1/2}$ ) for every water-Cherenkov detector
- **“Benchmark”**: Describe risetimes as a function of distance for a narrow energy range
- 1500 m array:  $19.1 \leq \log_{10}(E/\text{eV}) \leq 19.2$
- 750 m array:  $17.7 \leq \log_{10}(E/\text{eV}) \leq 17.8$
- Characterize every shower by a single parameter:  $\Delta_s$
- Golden-hybrid events (data-driven analysis):  $X_{\max}(\text{FD}) = f(\Delta_s)$
- Estimate  $X_{\max}(\text{SD})$



# Updated analysis

	<u>Published analysis</u> (Phys.Rev. D96 (2017) no.12, 122003)	<u>New analysis</u>
Data taking period	January, 1 2004 Dec, 31 2014	January, 1 2004 August, 31 2018
Energy range	0.3-100 EeV	3-100 EeV
Zenith angle range	< 45 degrees	< 60 degrees
Total number of events (1500m array)	58,583	125,005 (factor 2.1 bigger)
Events with $E > 20$ EeV	1,586	3,372 (factor 2.1 bigger)
Events with $E > 70$ EeV	49	106 (factor 2.1 bigger)



# Benchmark (1500 m array)

- Fit FADC traces from low-gain channel (A, B free parameters)

$$t_{1/2}^{\text{low-gain trace}} = 40\text{ns} + \sqrt{A(\theta)^2 + B(\theta)^2} - A(\theta)$$

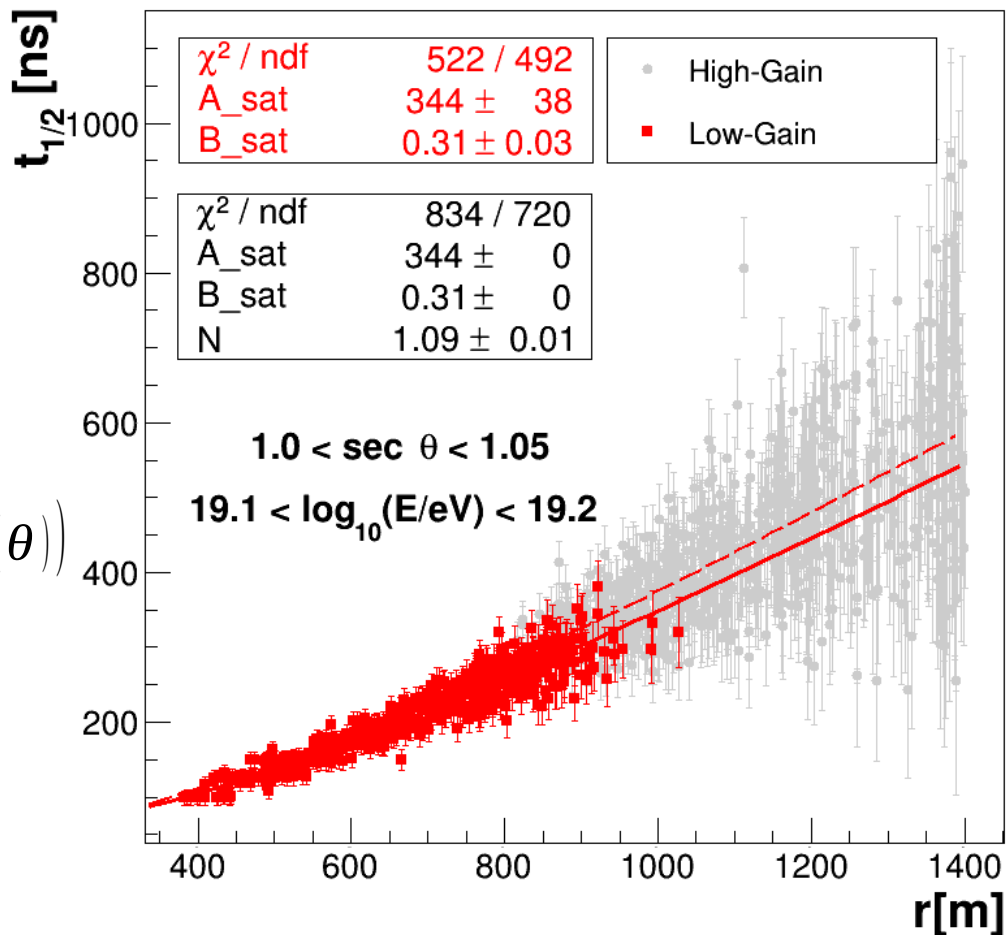
- Fit FADC traces from high-gain channel (only  $N$  is a free parameter)

$$t_{1/2}^{\text{high-gain trace}} = 40\text{ns} + N(\theta) \cdot \left( \sqrt{A(\theta)^2 + B(\theta)^2} - A(\theta) \right)$$

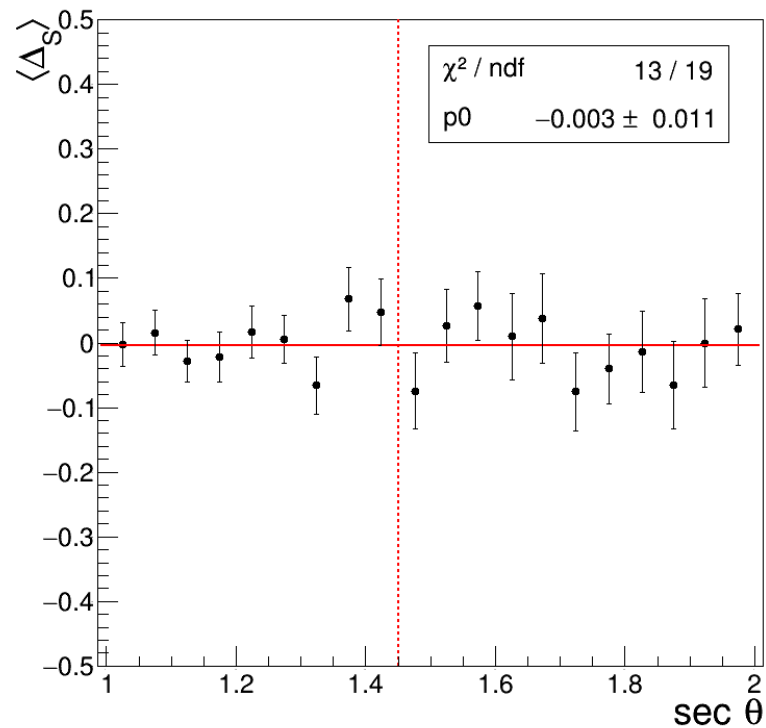
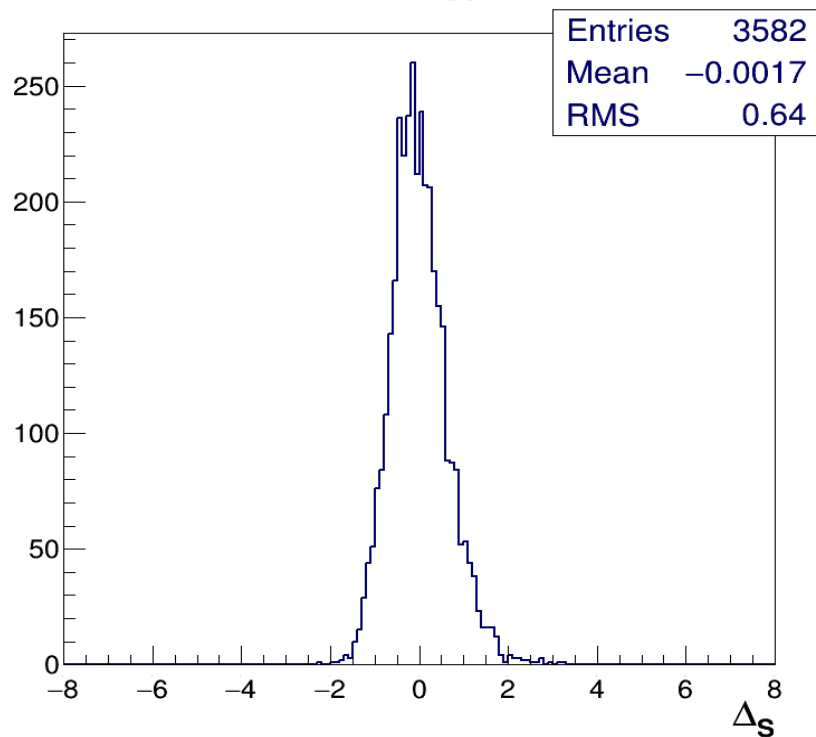
- Split parameterisations in two angular bins

$$\sec \theta < 1.45$$

$$1.45 < \sec \theta < 2.0$$



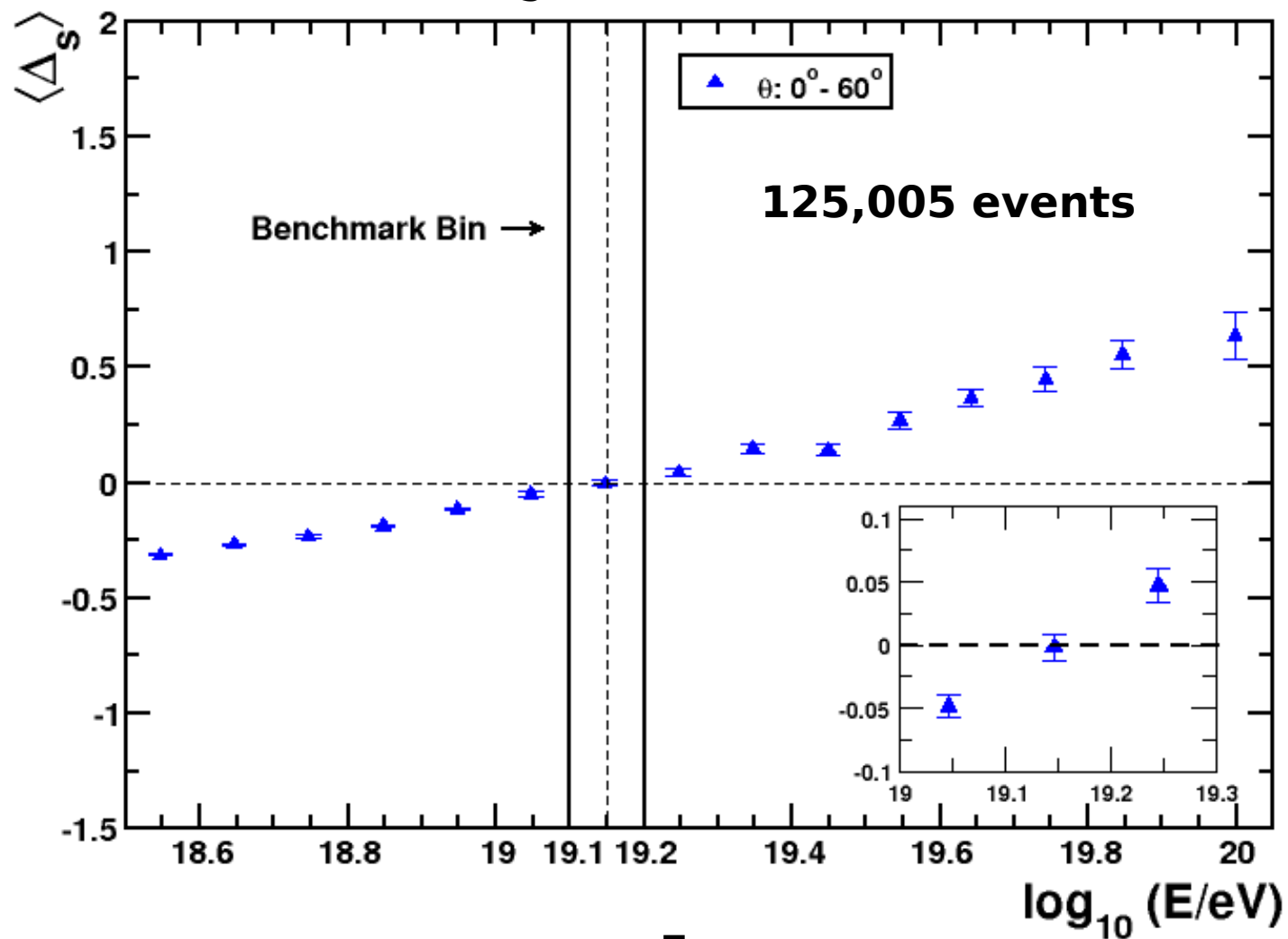
# $\Delta_s$ at benchmark energy



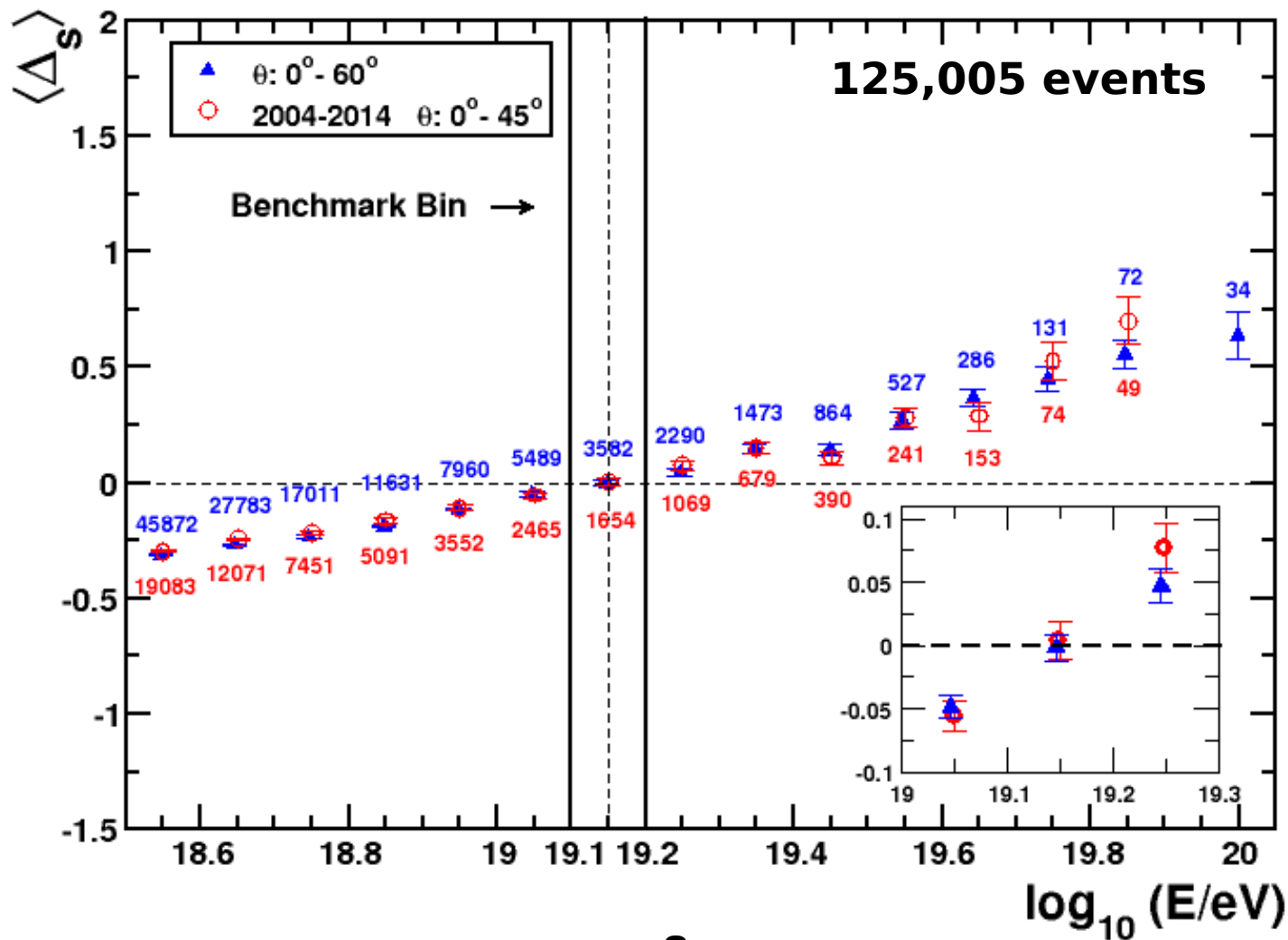
**$19.1 \leq \log_{10} (E/\text{eV}) \leq 19.2$**

**$\sec \theta \leq 2$**

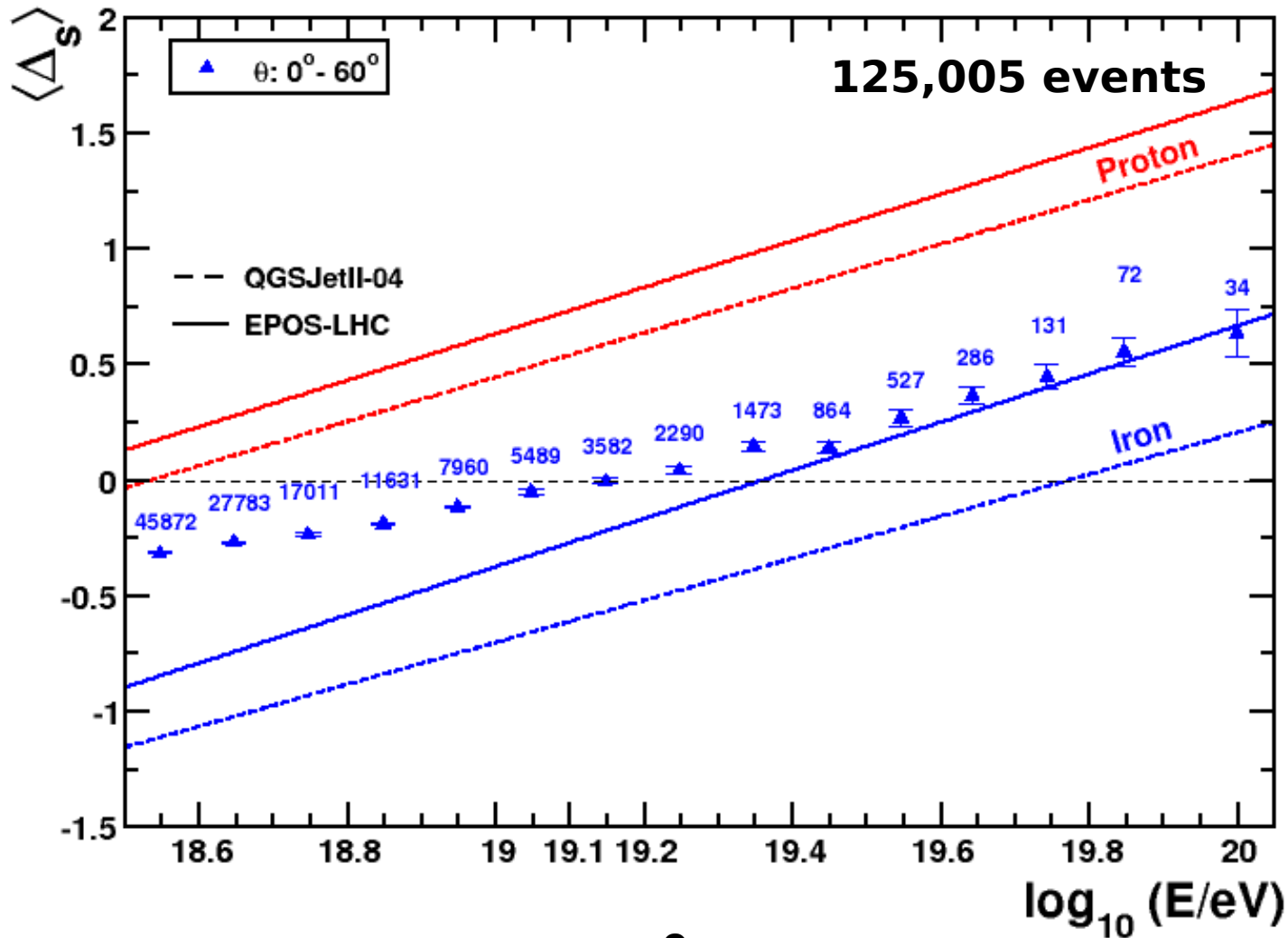
# $\langle \Delta_s \rangle$ vs Energy



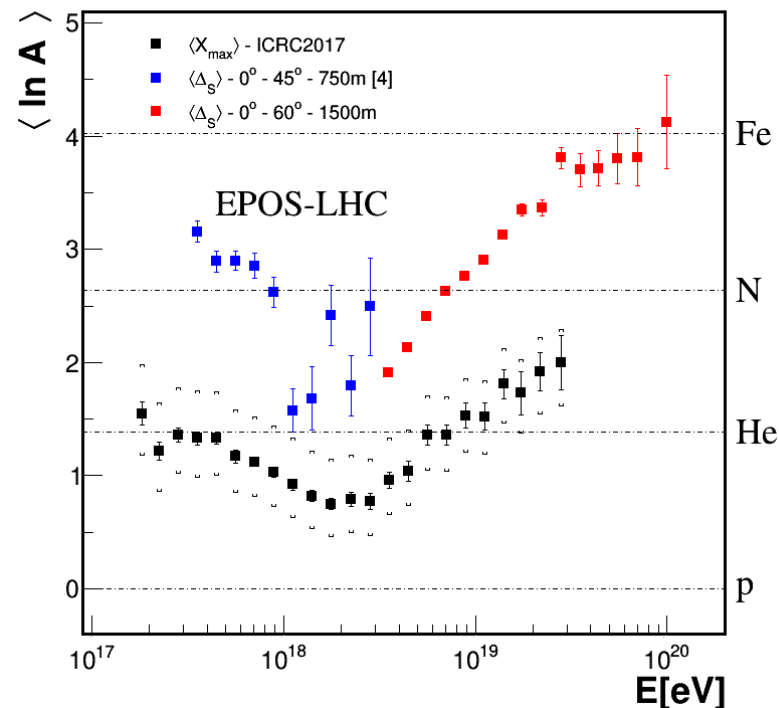
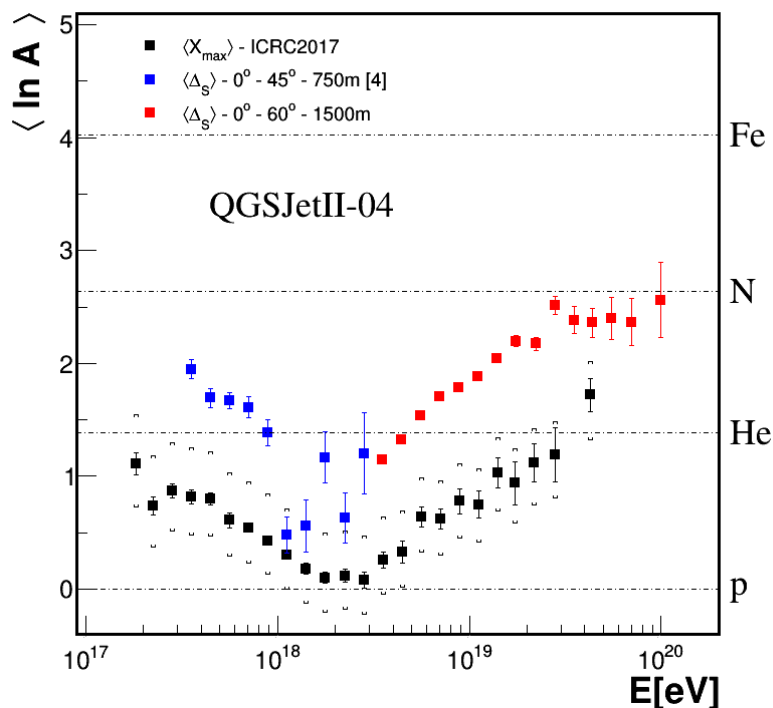
# $\langle\Delta_s\rangle$ vs Energy: Comparison with published analysis



# $\langle \Delta_s \rangle$ vs Energy: Comparison with models



# $\langle \ln A \rangle$ from $\langle \Delta_s \rangle$ compared to results from fluorescence measurements



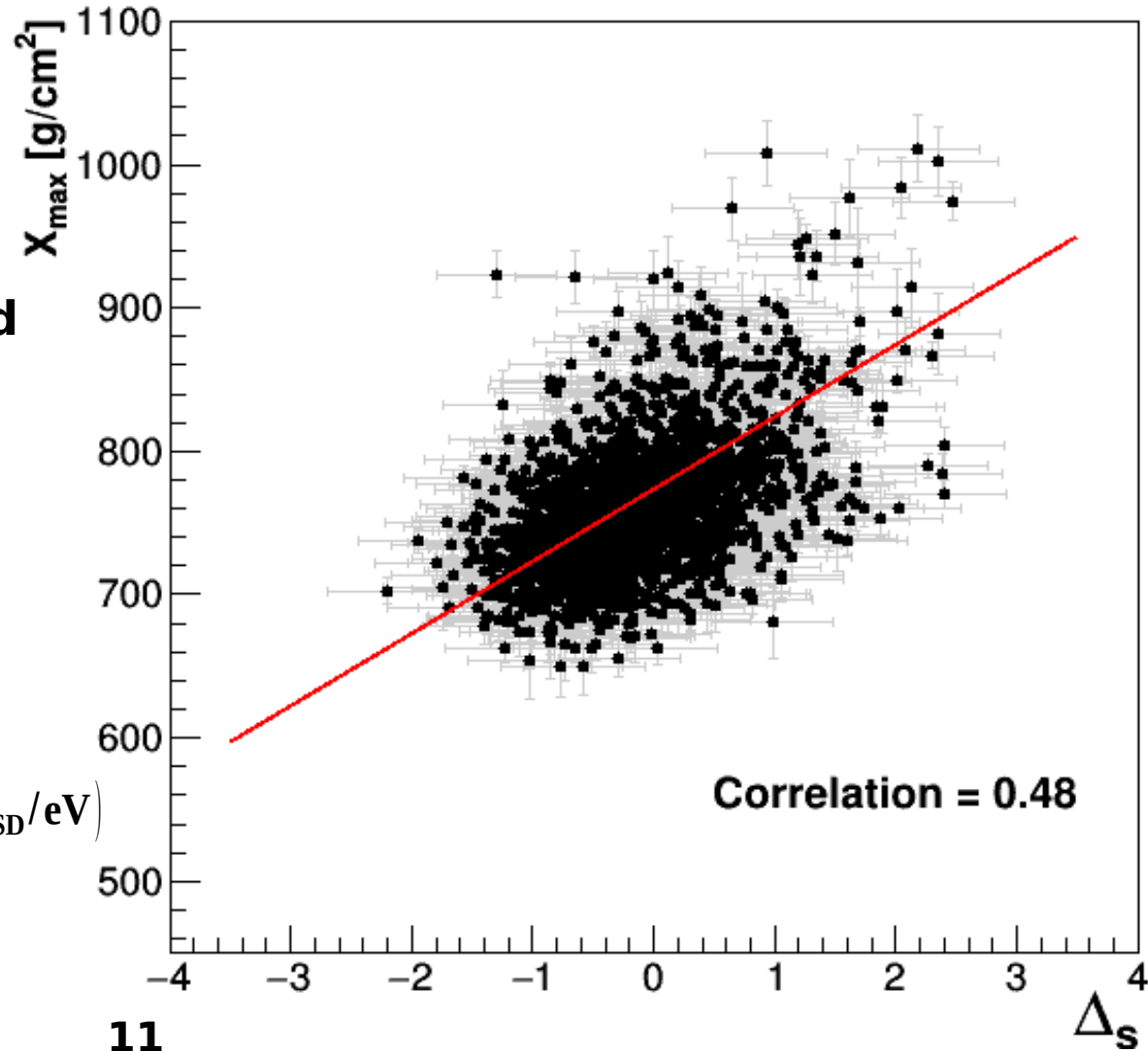
**Use models to convert measurements to  $\langle \ln A \rangle$**

- Simulations do not describe properly the signals recorded by Surface Detectors [Muon Problem addressed in PoS(ICRC2019)214,404,411]
- To estimate mass composition, data from Surface Detectors must be calibrated with fluorescence measurements

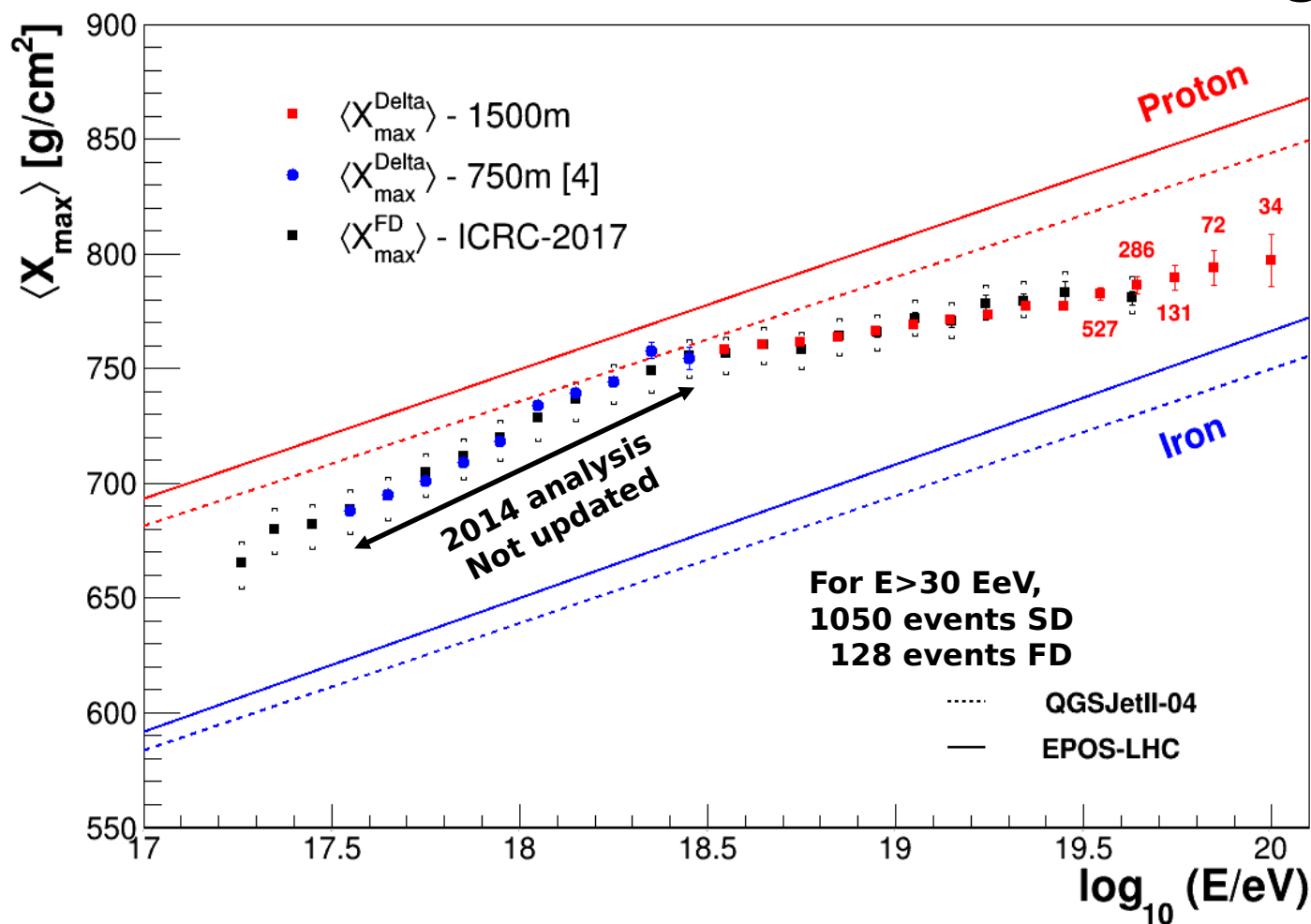
# Calibration with $X_{\max}$

- Use golden hybrids (events simultaneously reconstructed by SD and FD)
- 1500 m array : **2126 events**;  
Correlation factor: 0.48

$$\langle X_{\max} \rangle = (733 \pm 38) + (34 \pm 12) \cdot \langle \Delta_s \rangle + (2 \pm 1) \cdot \log_{10}(E_{\text{SD}}/\text{eV})$$



# Evolution of $\langle X_{\max} \rangle^{\text{SD}}$ as a function of energy



Preponderance of intermediate/heavy nuclei as energy increases



# Thank you for your attention



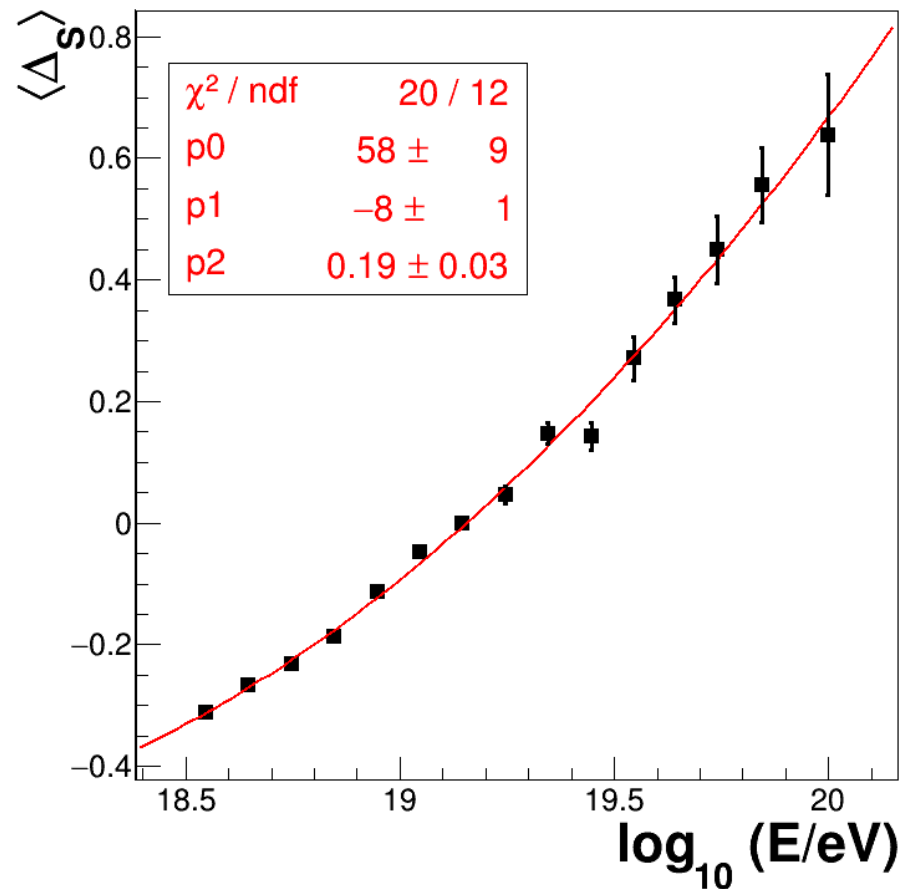
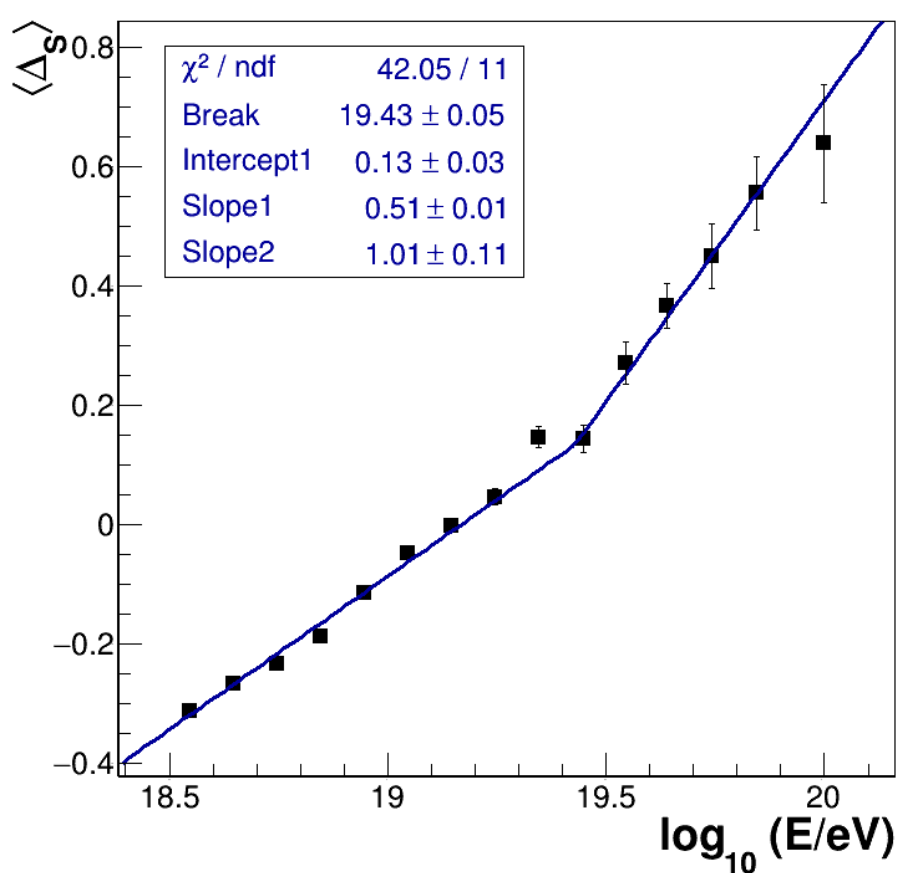
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# Backup



**The rate of change of mass composition varies with energy**